

CLUSTER



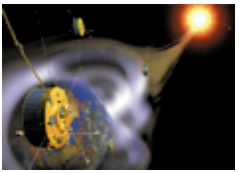
esa



# Cluster–Double Star–THEMIS Science Coordination

*C.P. Escoubet<sup>1</sup>, M. Hapgood<sup>2</sup>, M.L. Goldstein<sup>3</sup>, and M.G.G.T. Taylor<sup>1</sup>,*  
*<sup>1</sup>ESA/ESTEC (NL), <sup>2</sup>RAL (UK), <sup>3</sup>GSFC (USA),*

- Some Cluster results (relevant to THEMIS objectives)
- Cluster–Double Star–THEMIS coordination
  - a few orbital conjunctions



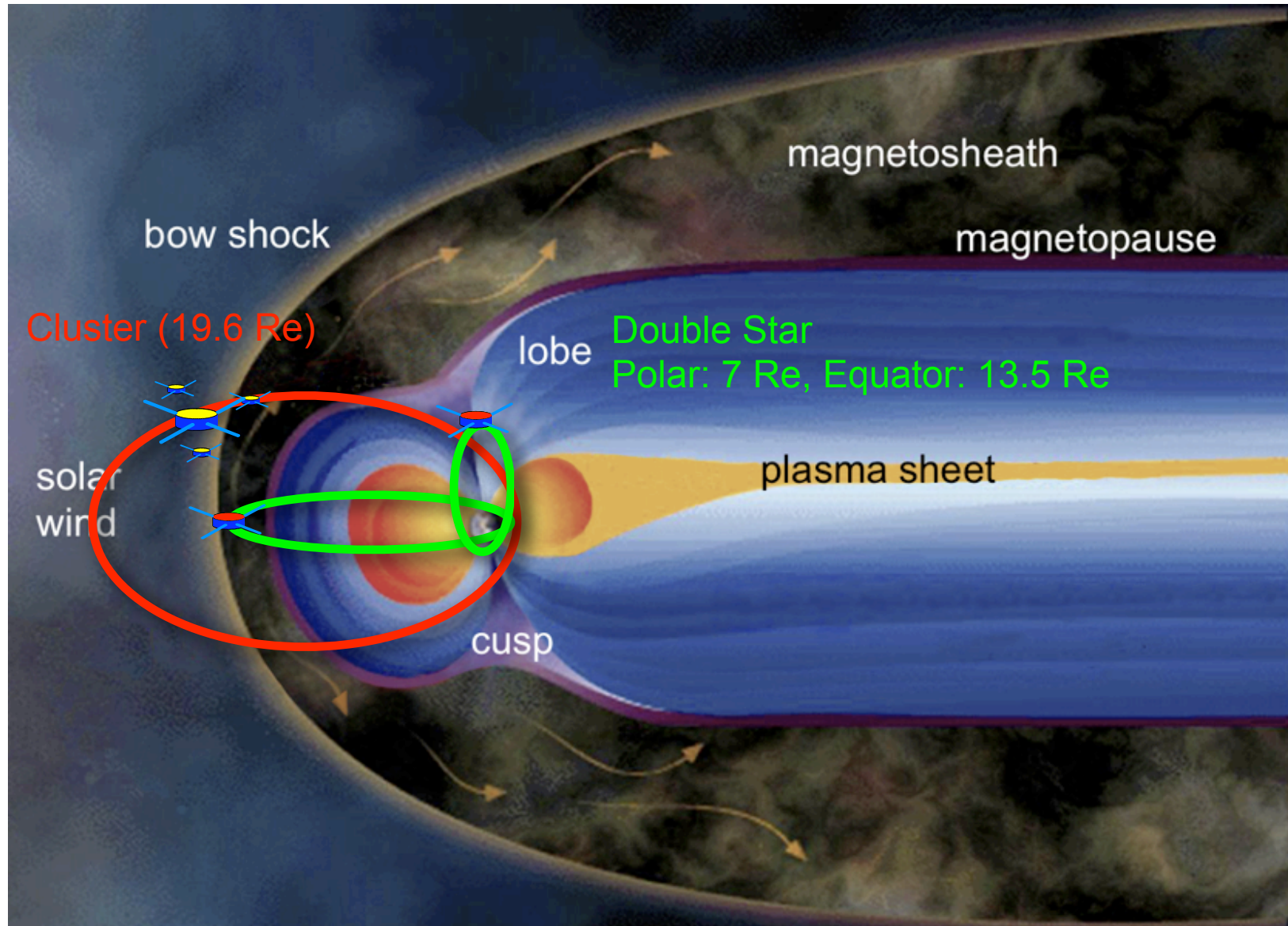
CLUSTER

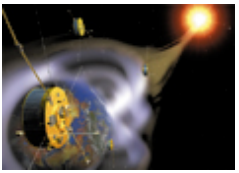


esa



# Cluster–Double Star orbits, dayside

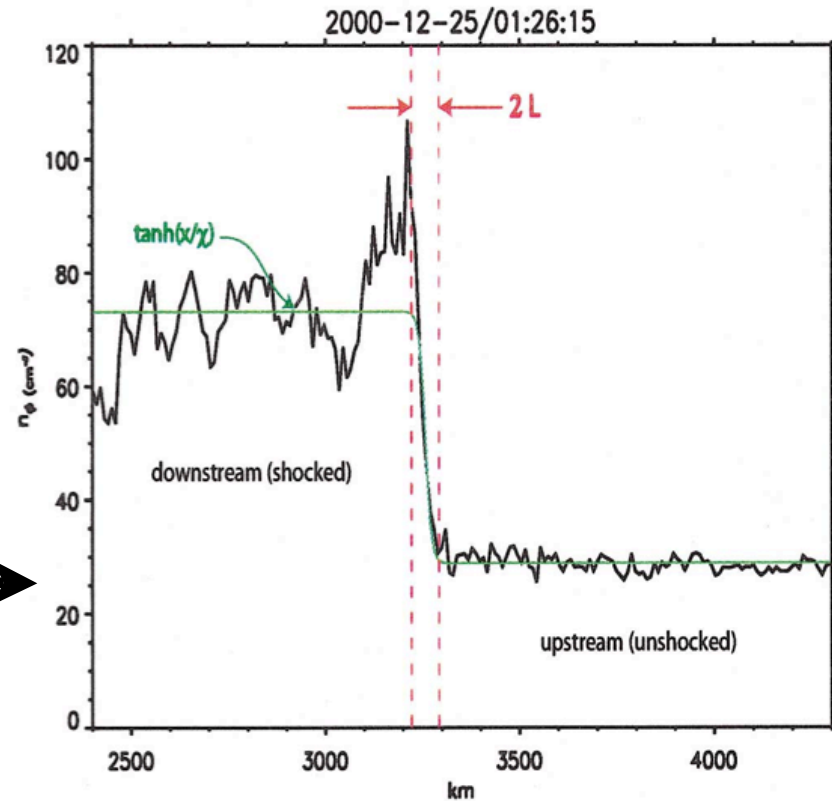
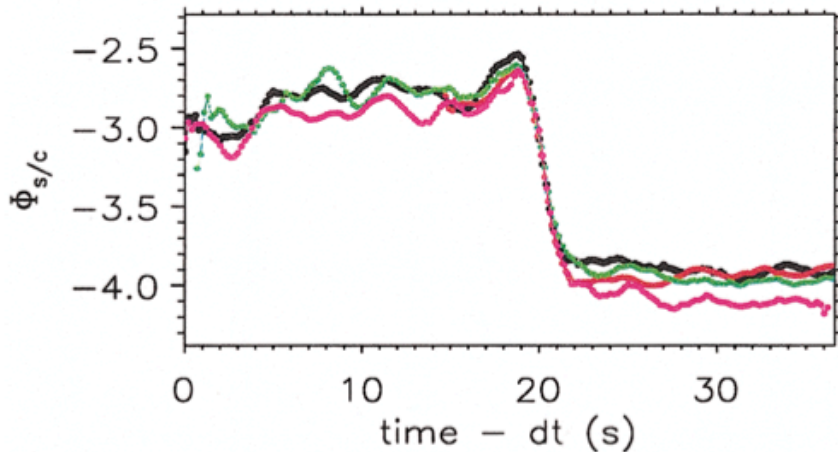
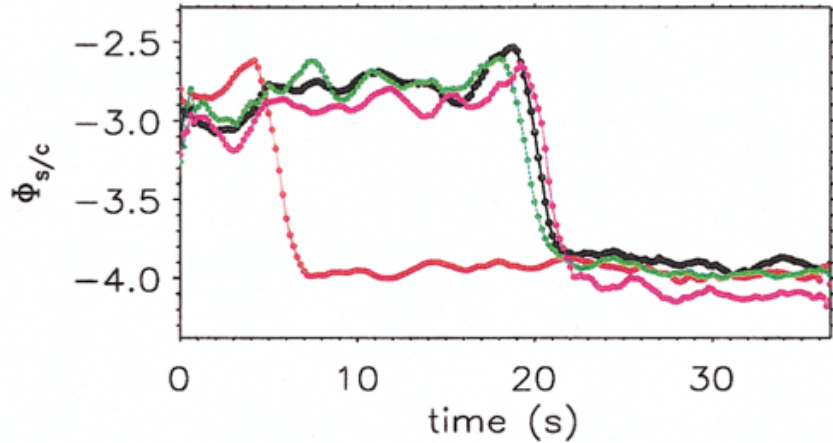




CLUSTER

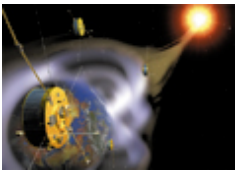


# Bow shock thickness using 4 s/c



- bow shock thickness proportional to downstream ion gyroradius

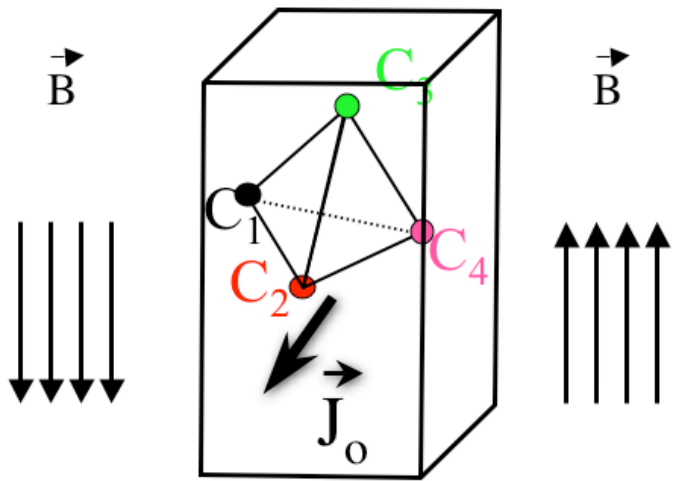
[Bale et al., 2003]



CLUSTER

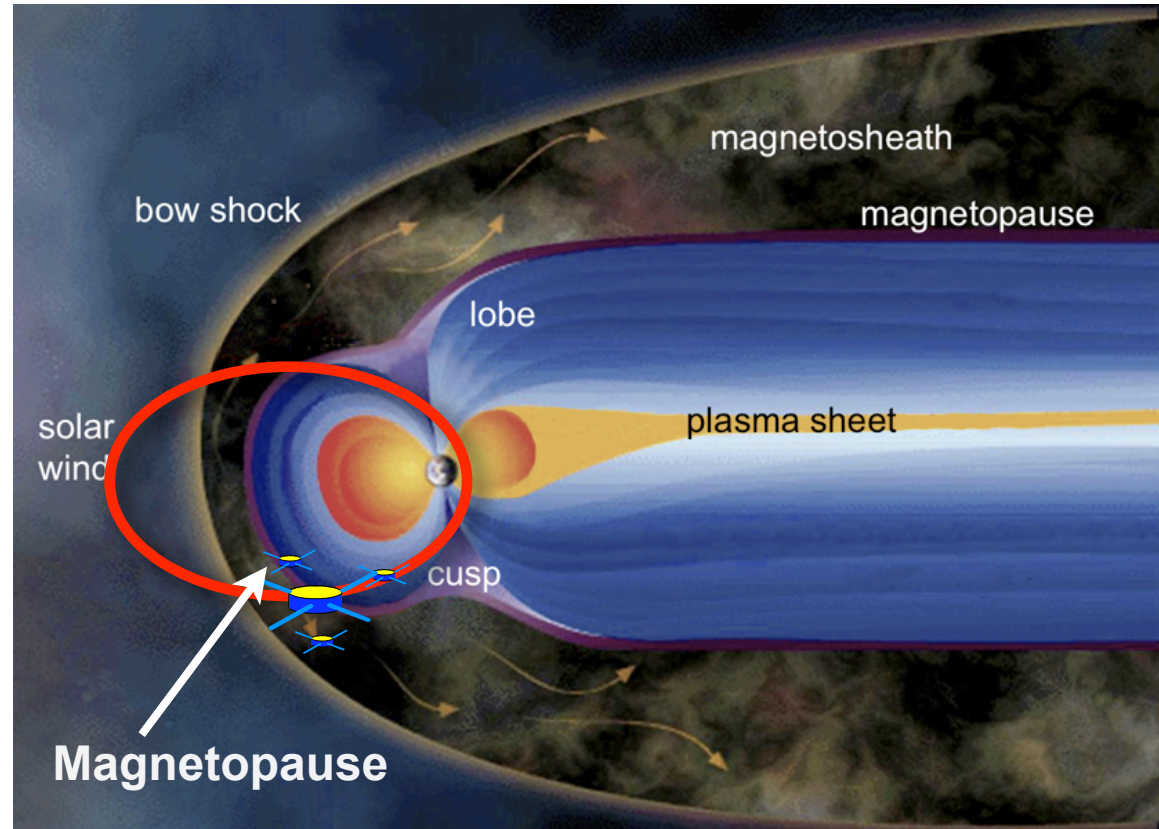


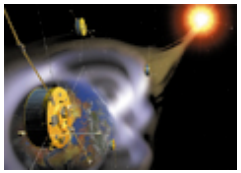
# Chapman-Ferraro Current



Solar Wind | Magnetopause | Magnetosphere

Current is given by 
$$\vec{J}_0 = \text{curl } \vec{B}$$





CLUSTER

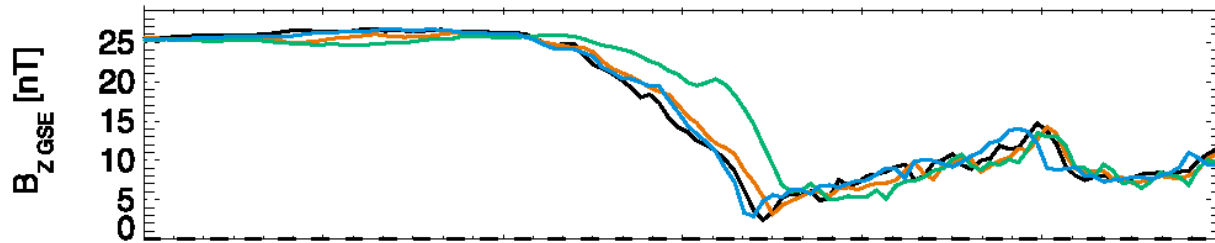


esa

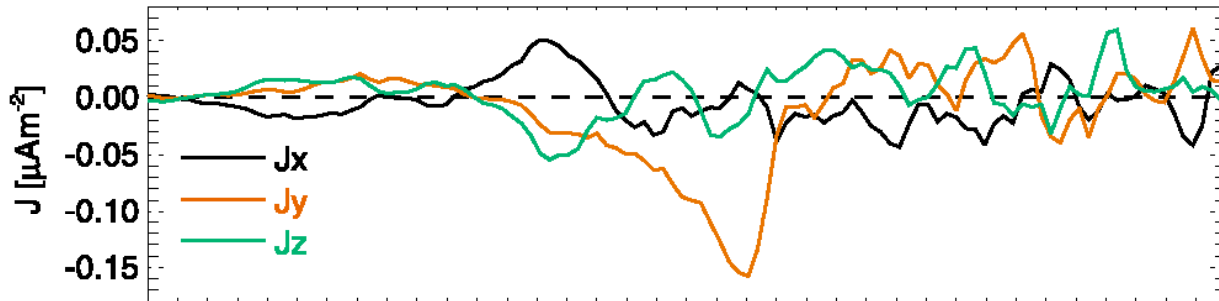


# Magnetopause current density with 4 s/c

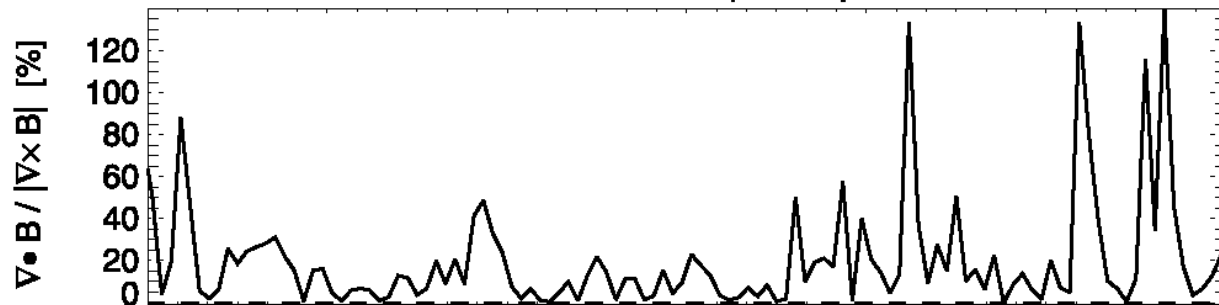
Magnetopause: 200 km



Current density from  $\nabla \times B$



$\nabla \cdot B / |\nabla \times B|$

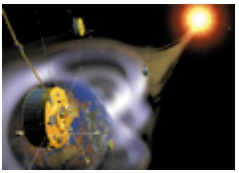


UT 0330:51 0330:57 0331:03 0331:10 0331:16 0331:22 0331:29

- Magnetopause Thickness and current

- s/c separation 100 km
- all s/c inside current sheet
- $J=0.15 \text{ uAm}^{-2}$

[Haaland et al., 2004]



CLUSTER

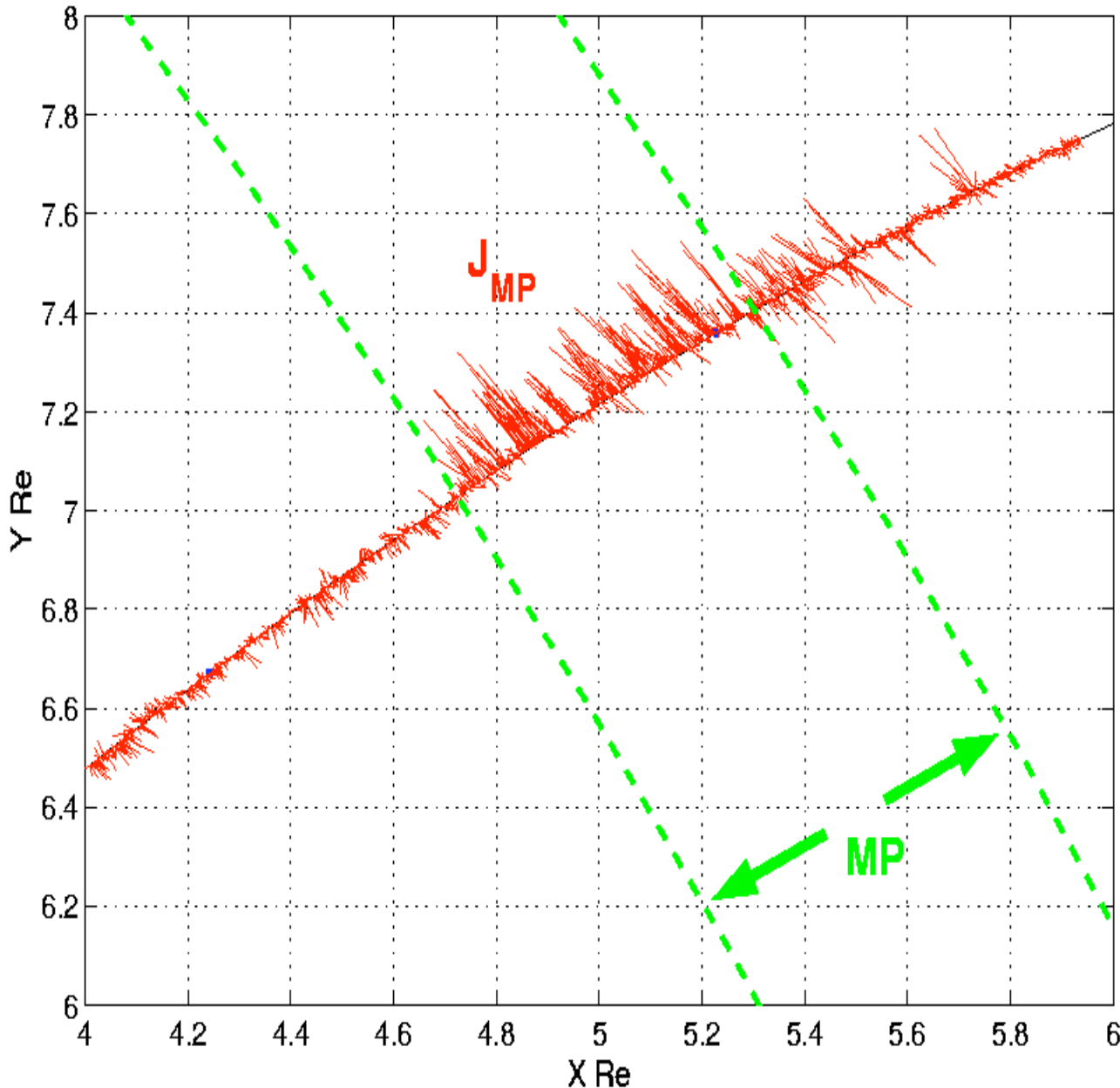


esa



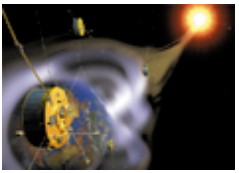
Cluster - Orbit & s/c - configuration

2001 Jan 26



- Current vector plotted along s/c trajectory
- Magnetopause current variability due to magnetopause motion
- Current vector aligned with magnetopause model

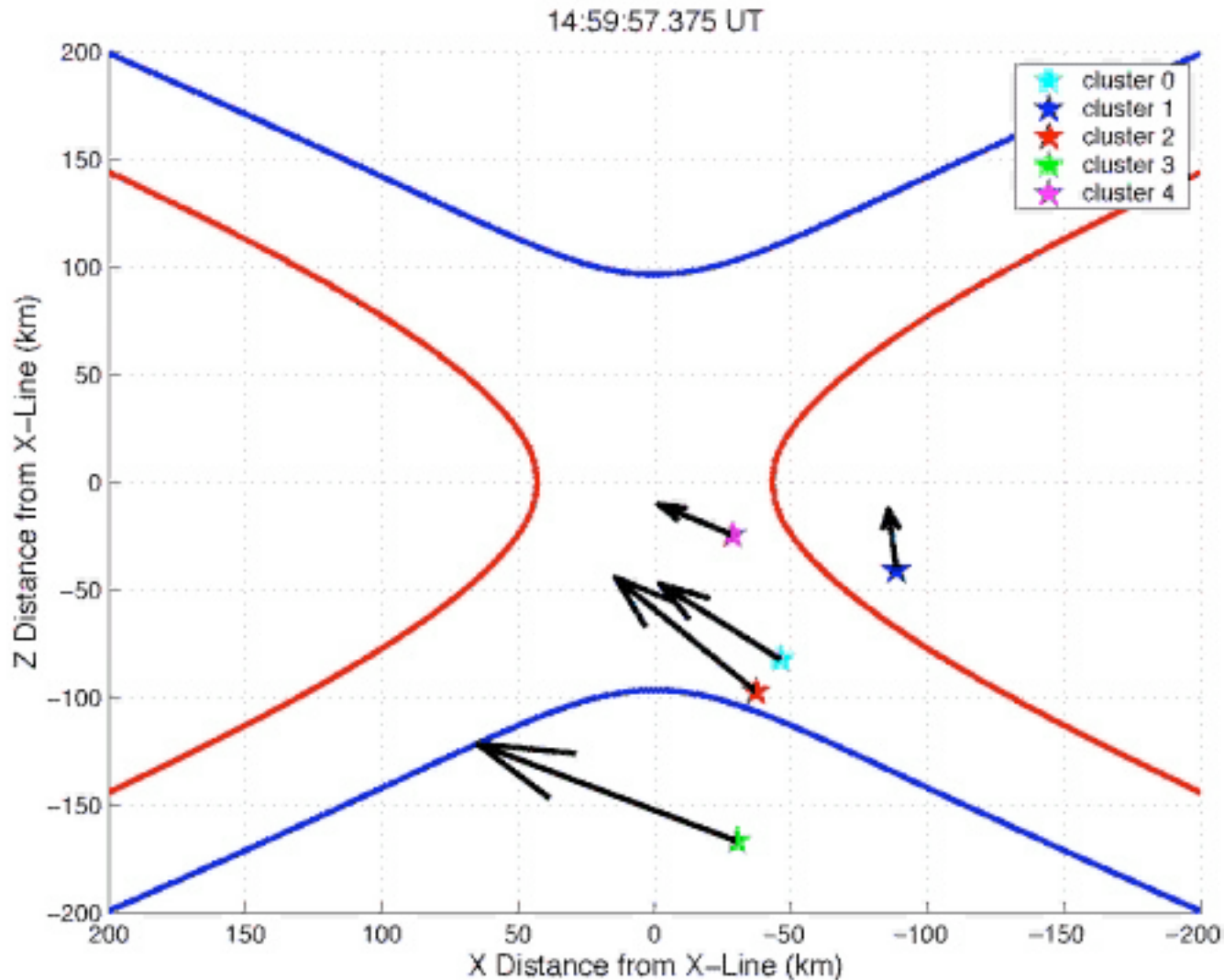
[Dunlop et al., 2002]



CLUSTER



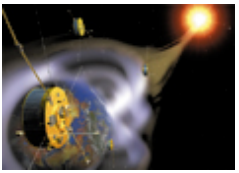
# Reconnection observed in B direction



18 March 2002

- Spacecraft separation: 100km
- Cluster 2 (red): closest approach at 1 km from X line
- Cluster 4: 3.5 km from X line

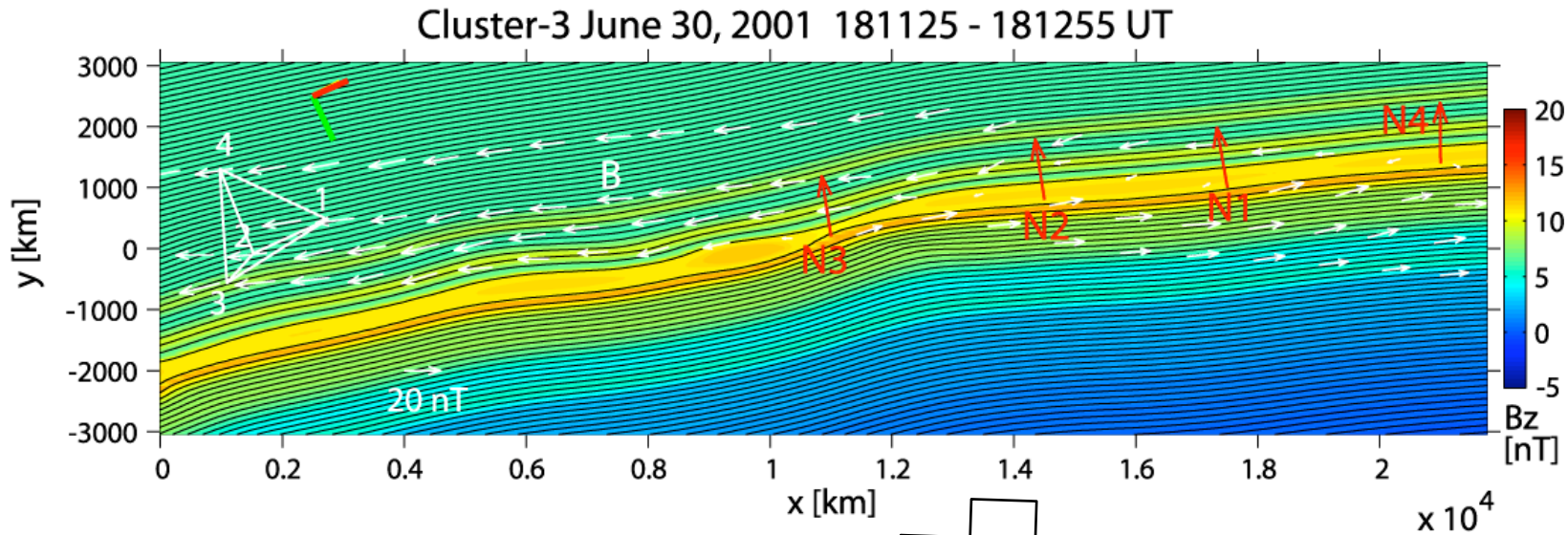
[Wendel and Reiff, 2006]



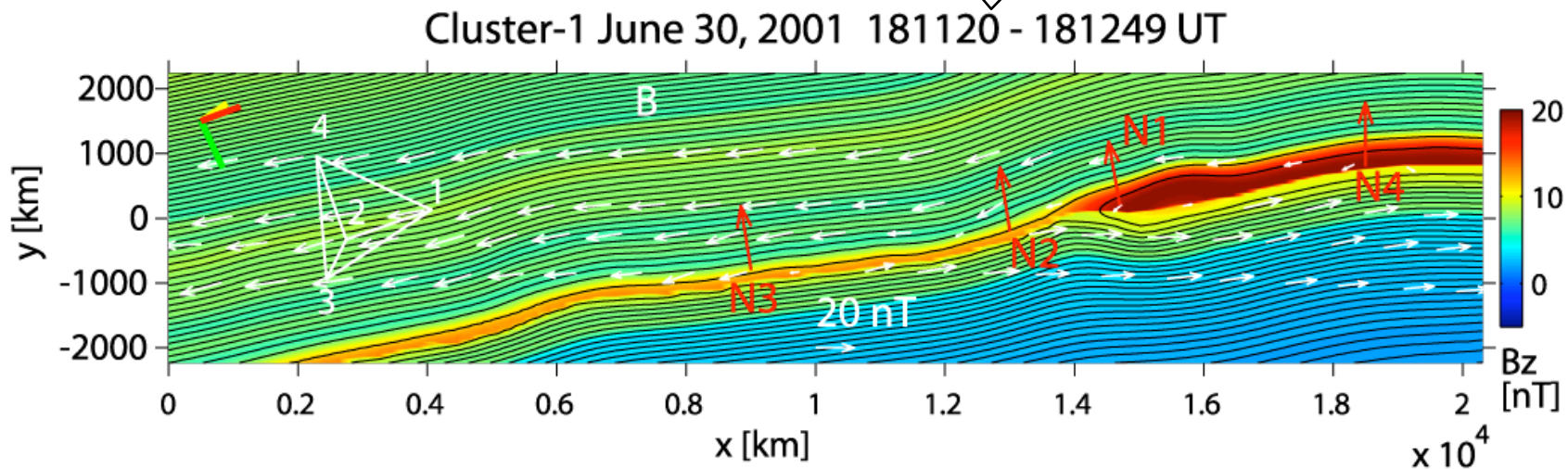
CLUSTER



esa

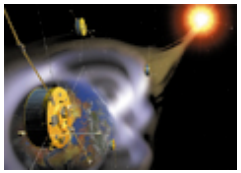


Reconnection onset



[Hasegawa and Sonnerup, 2003]



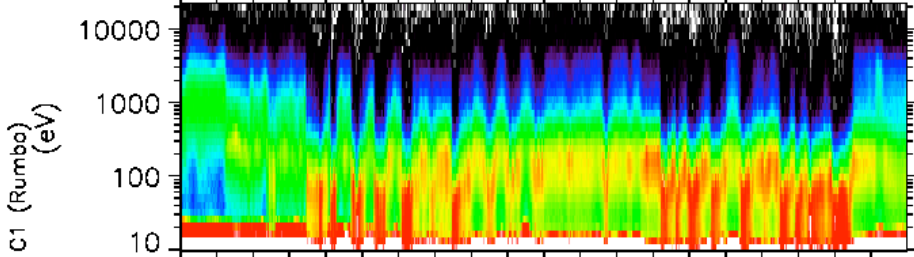


CLUSTER

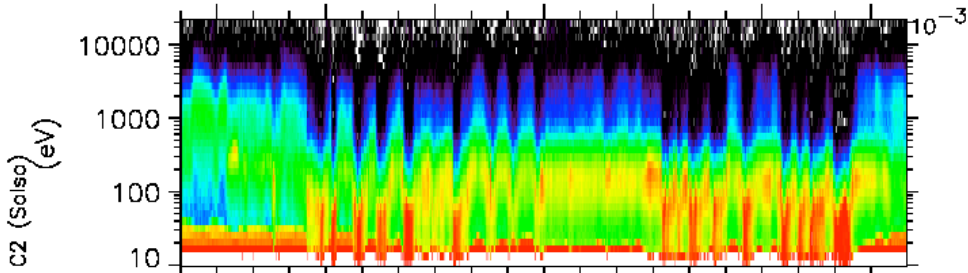


14/06/2001 (165) Energy, pitch angles av. 0–180deg.

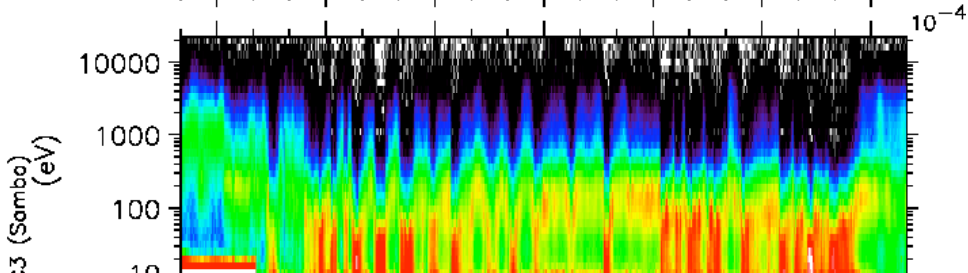
C1



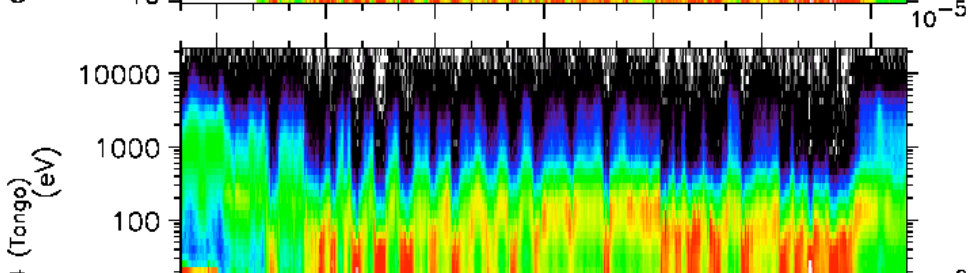
C2



C3

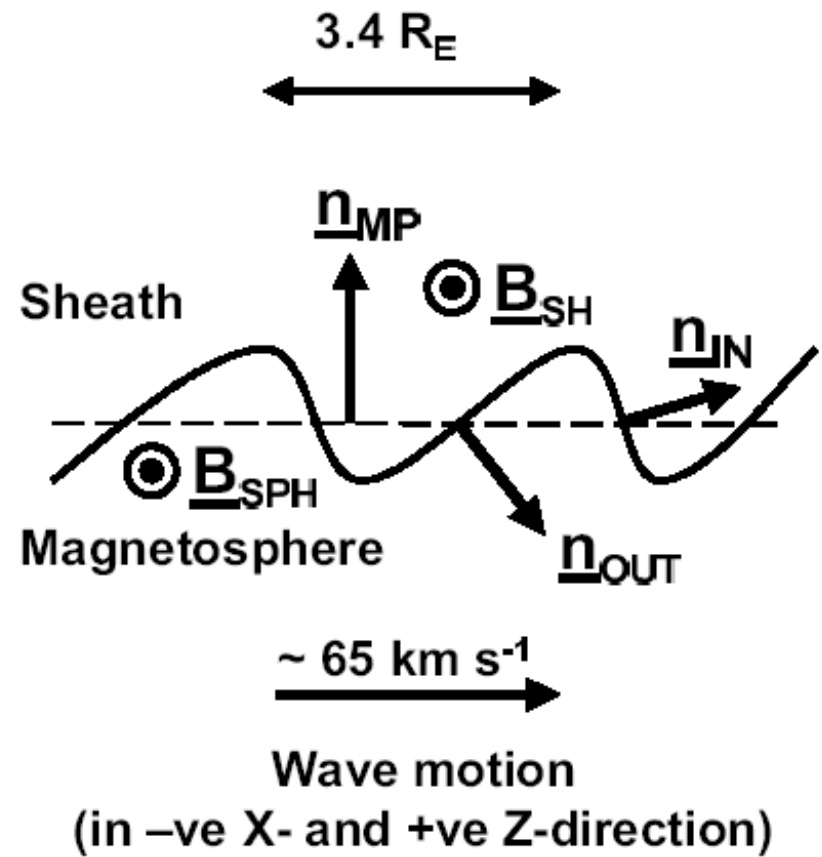


C4

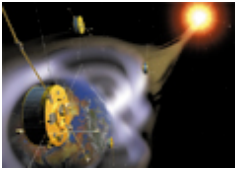


UT	15.45	16.00	16.15	16.30	16.45	17.00	17.15	hr:min
X GSE	-5.559	-5.583	-5.606	-5.628	-5.648	-5.667	-5.685	$R_E$
Y GSE	-16.696	-16.569	-16.439	-16.305	-16.167	-16.026	-15.881	$R_E$
Z GSE	-4.024	-4.153	-4.280	-4.407	-4.532	-4.656	-4.780	$R_E$

## Kelvin-Helmholtz waves



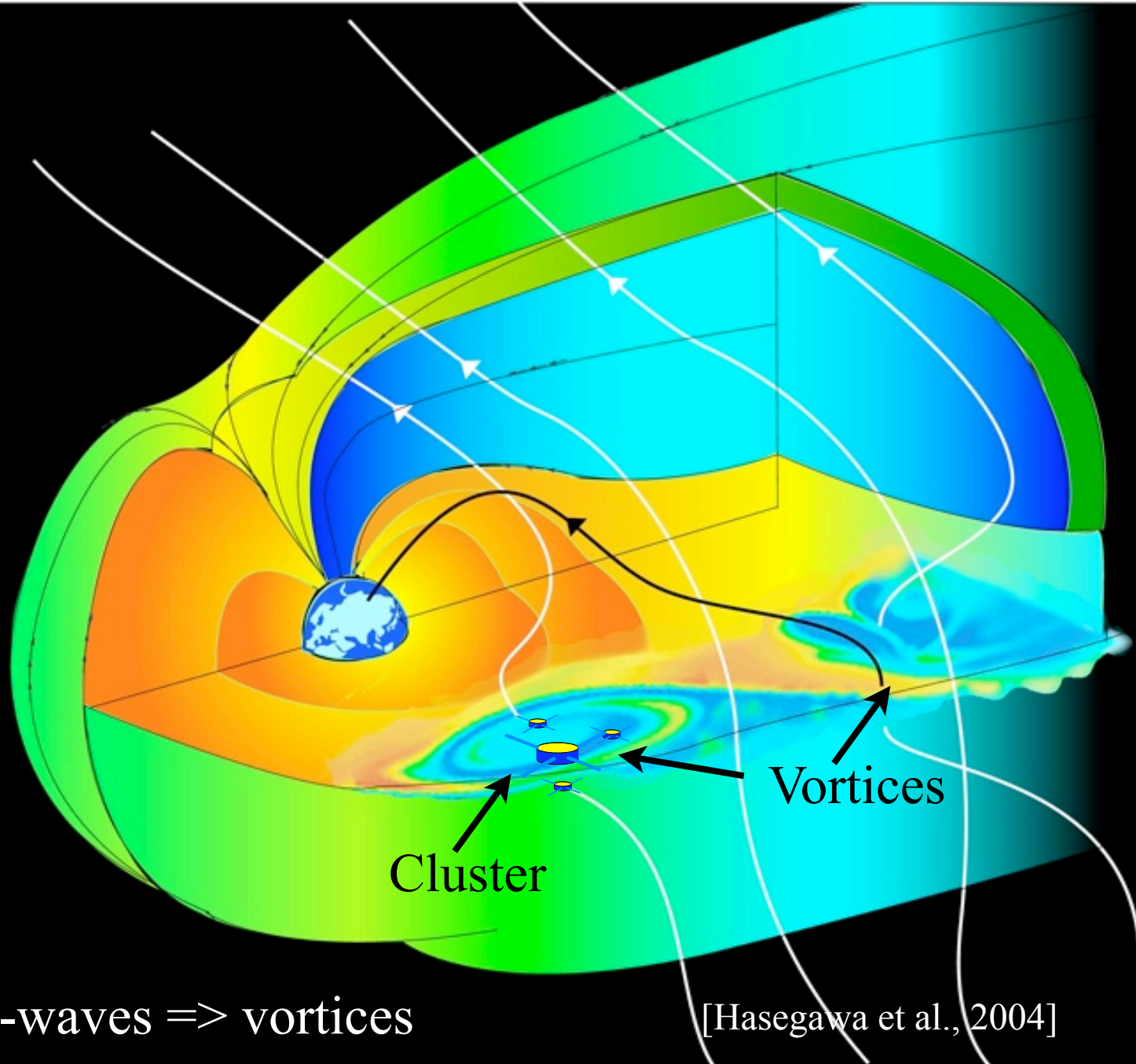
[Owen et al., 2003]



CLUSTER

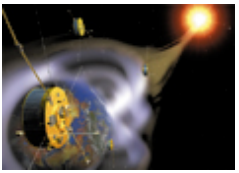


esa

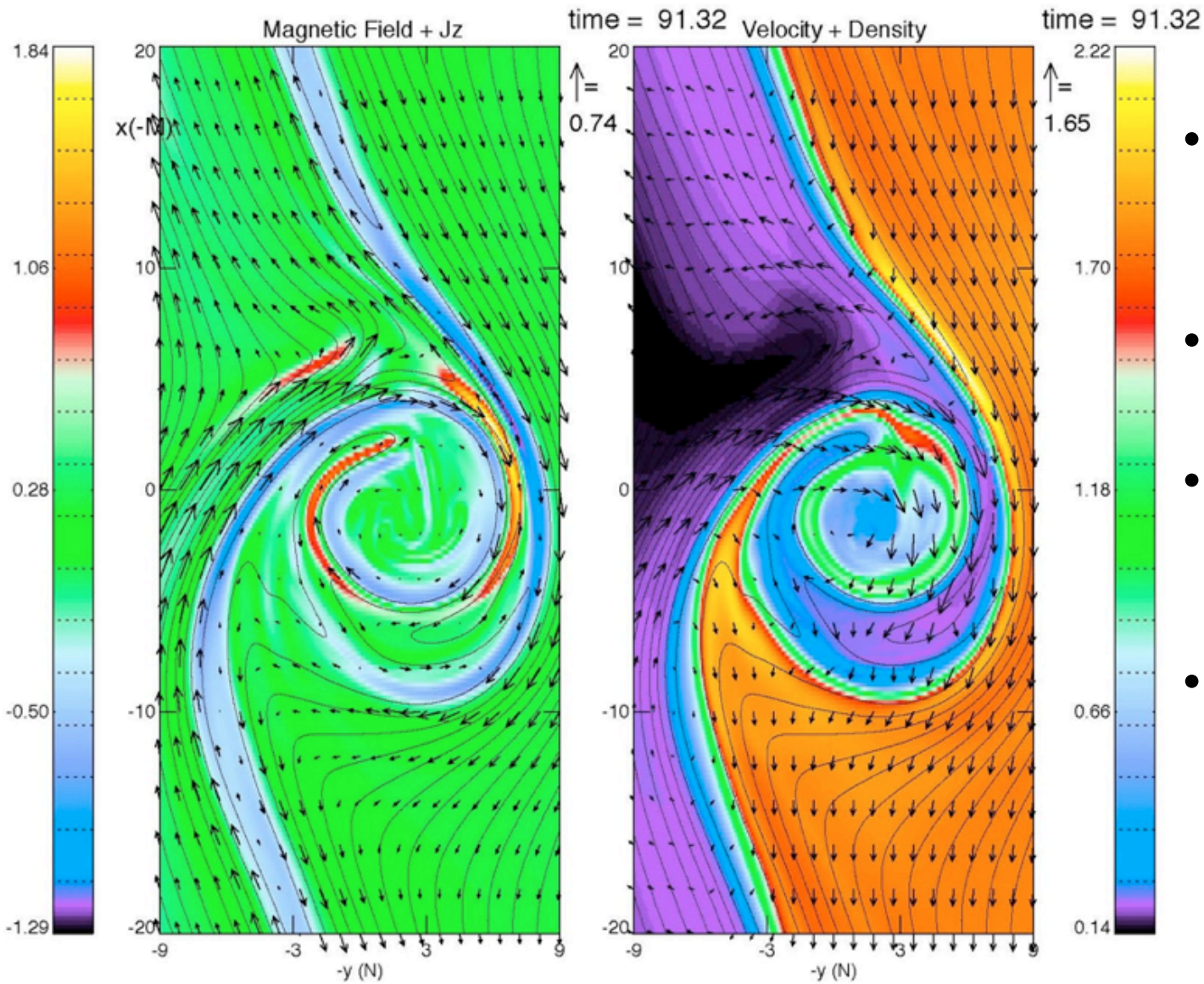


Rollup-waves => vortices

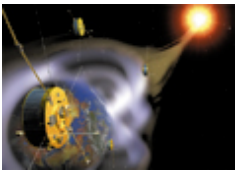
[Hasegawa et al., 2004]



CLUSTER



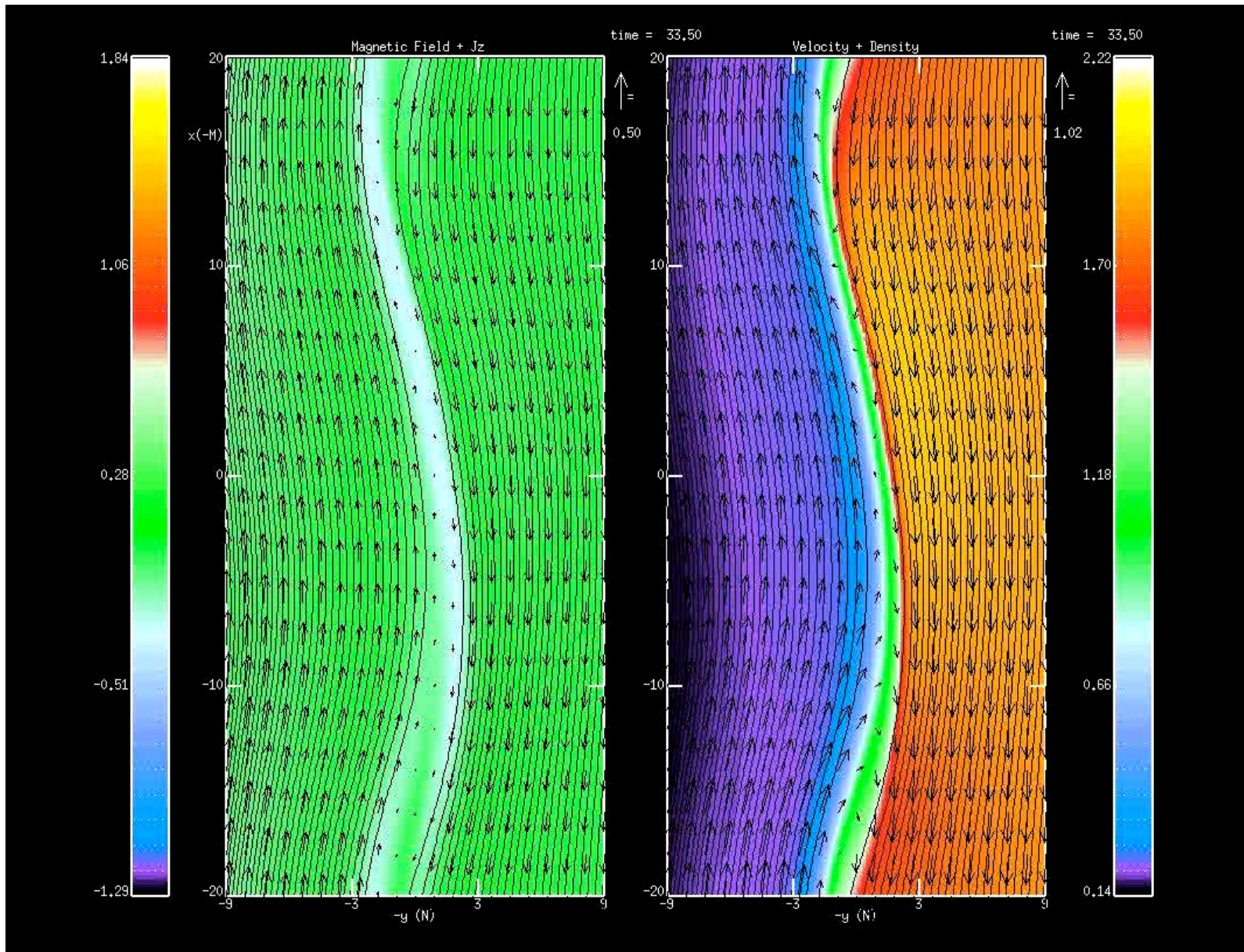
- MHD simulation of Kelvin-Helmholtz vortices
- Reconnection taking place inside vortices
- Also observed on Cluster CIS data (Walen test)
- wave length  $\sim 6Re$

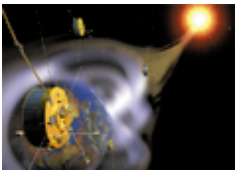


CLUSTER



esa





CLUSTER

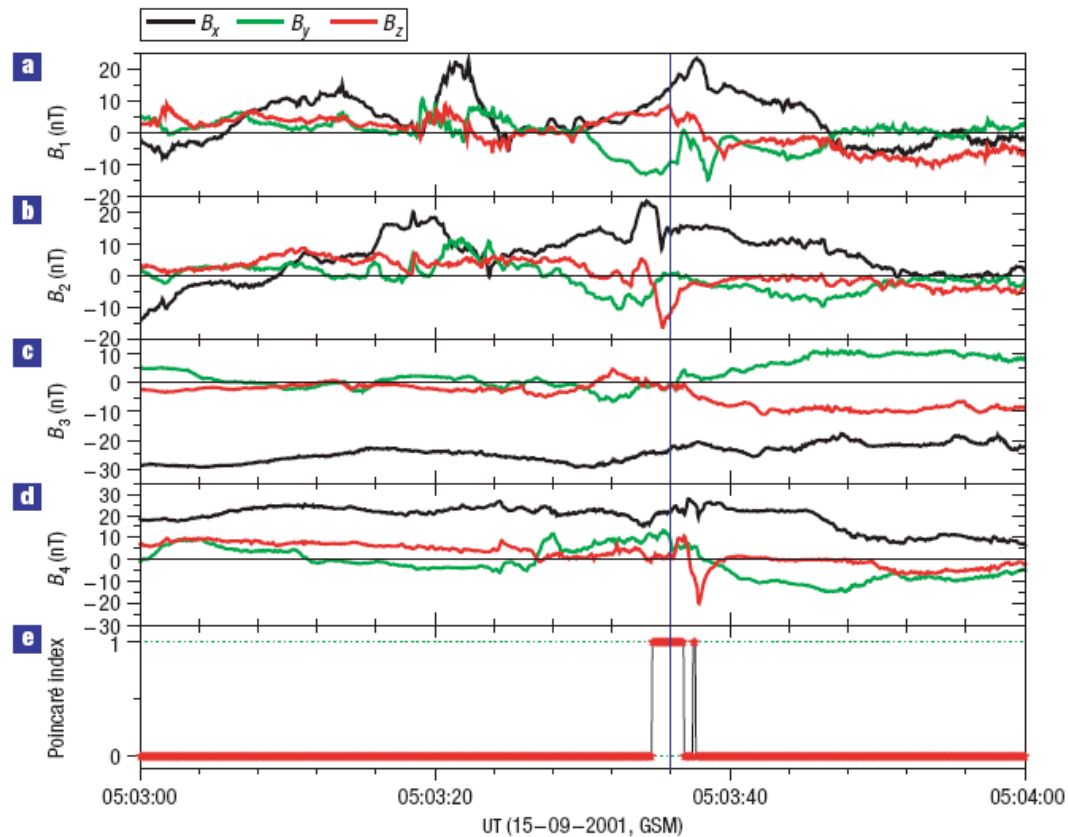


esa

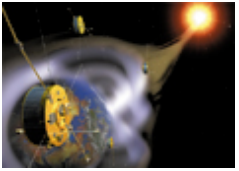


# Magnetic Null

- (C.J. Xiao *et al.*, Nature Physics, 2 July 2006)



**Figure 3** The high-resolution (0.04 s) magnetic field data of four Cluster spacecraft, and the calculated Poincaré index during 05:03–05:04 UT on 15 September 2001. a–d, The magnetic field data of C1–C4. The black, green and red lines indicate the x, y and z component, respectively. e, The Poincaré index calculated from intercalibrated high resolution (0.04 s) data. The vertical blue line shows the time 05:03:36 UT when a null point is found to exist inside the Cluster tetrahedron based on 4-s resolution data.



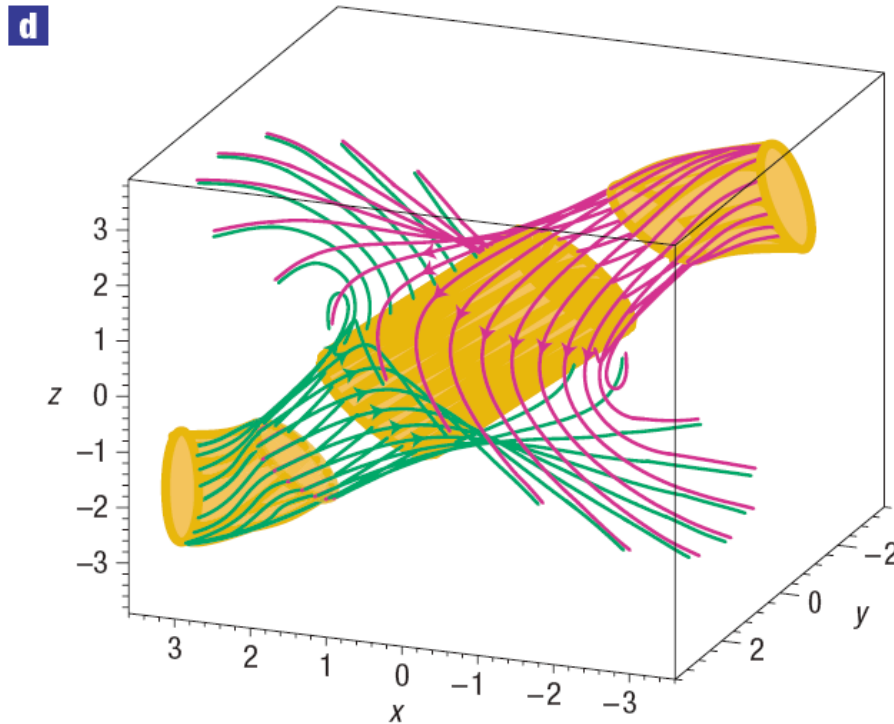
CLUSTER



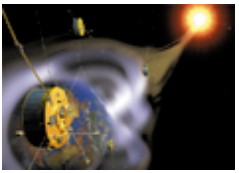
esa



# Magnetic Null ...



The inferred spiral structure of the magnetic field  
around the null



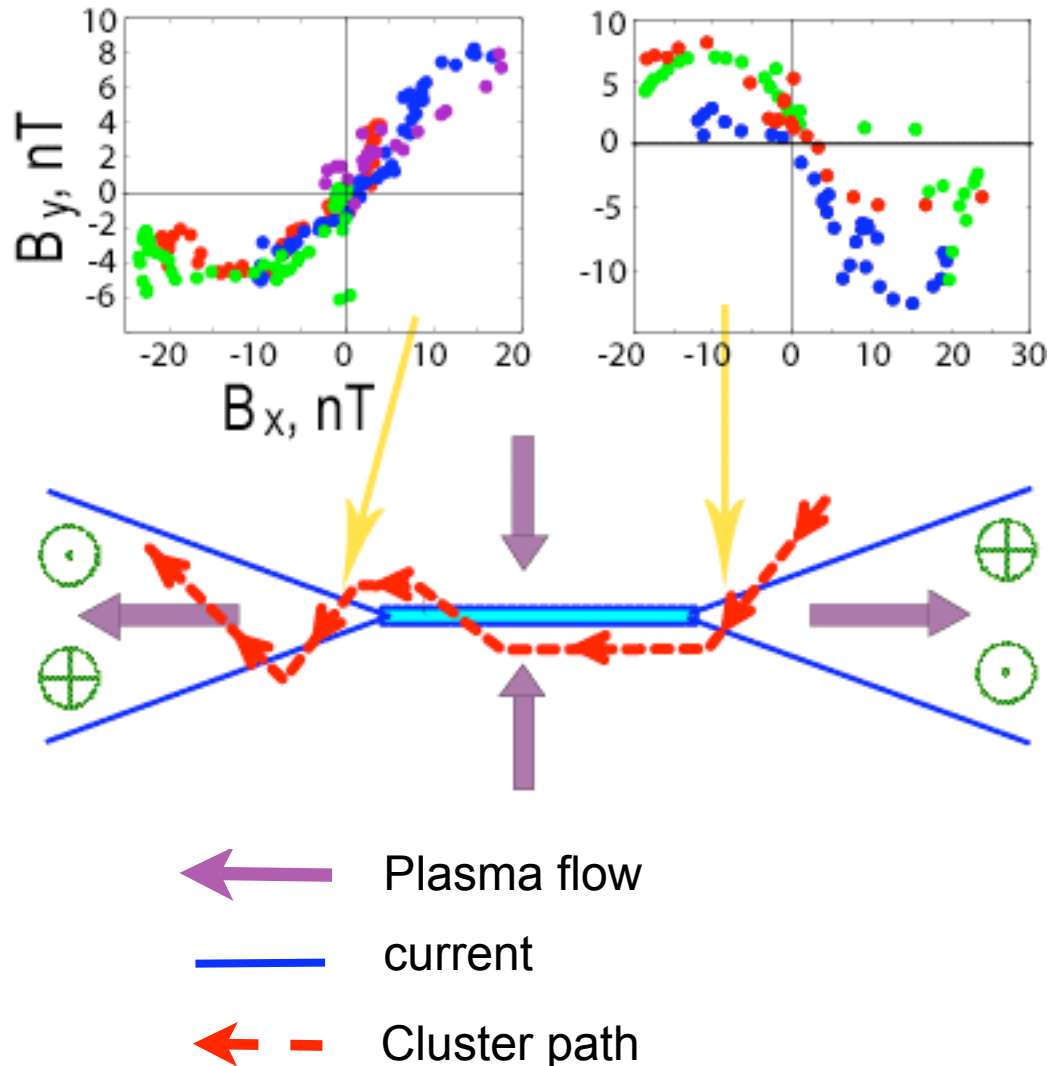
CLUSTER



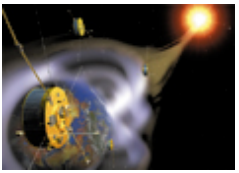
esa



## Bifurcated current sheet



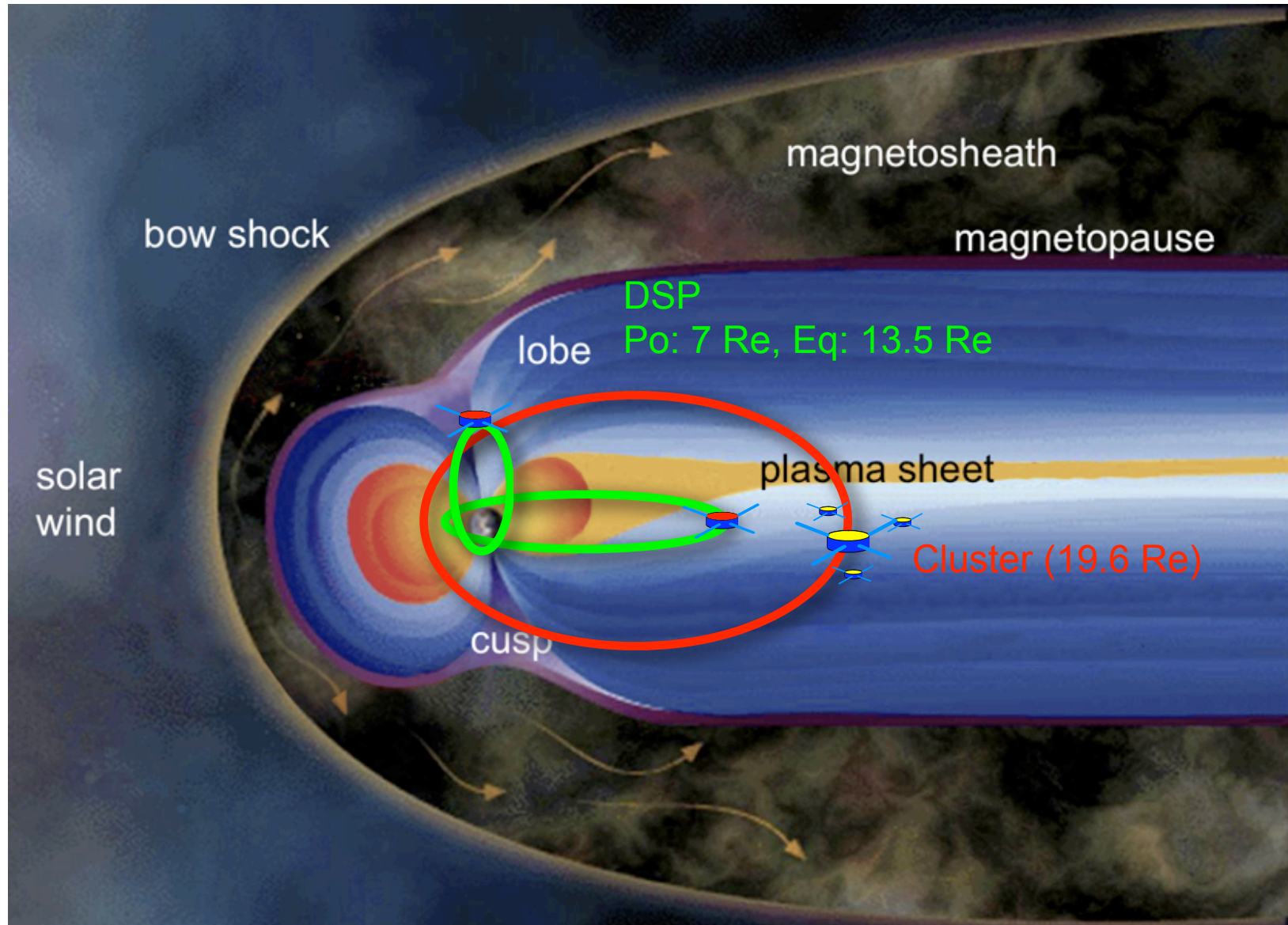
- Hall Effect ( $\delta B_y$ ) measured during outer crossing
- 500 km thin current sheet around an X-line bifurcated on both sides
- Flow reversal and field line curvature reversal are proof that an X-line moved tailward over Cluster



CLUSTER



# Cluster–Double Star orbits, tail



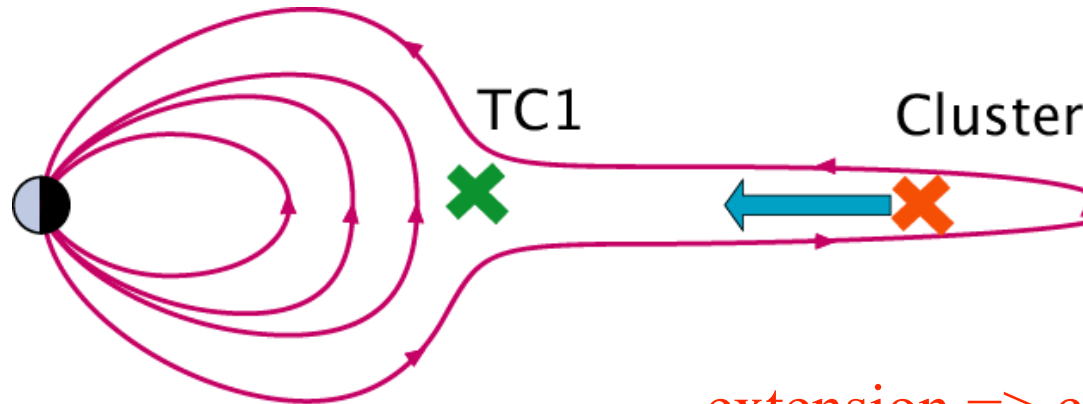




# Double Star

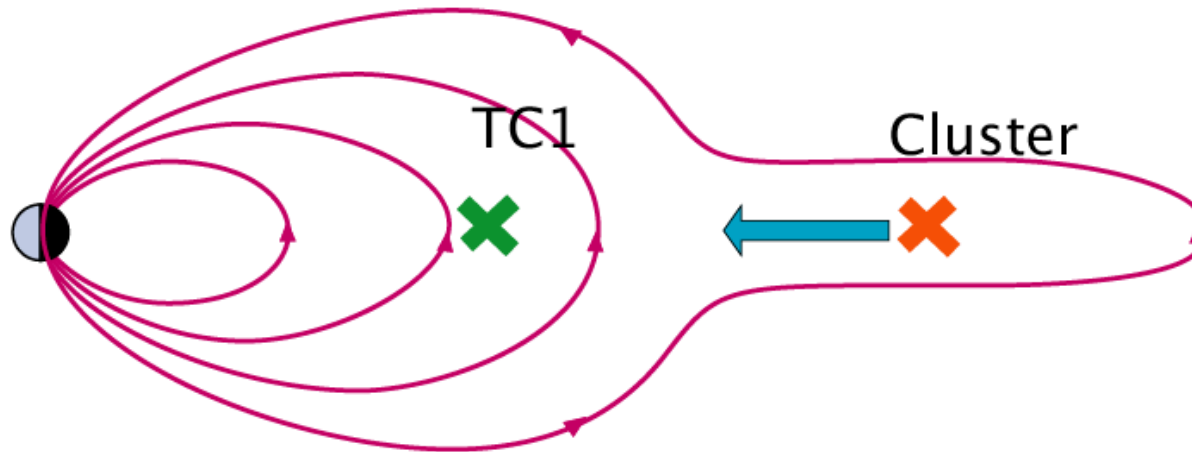


Bursty bulk flow and dipolarization

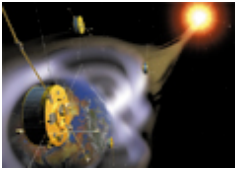


extension => comparison midnight-morning

Bursty bulk flow but no dipolarization



→ The magnetic topology in the inner magnetosphere can control the earthward penetration of dipolarization

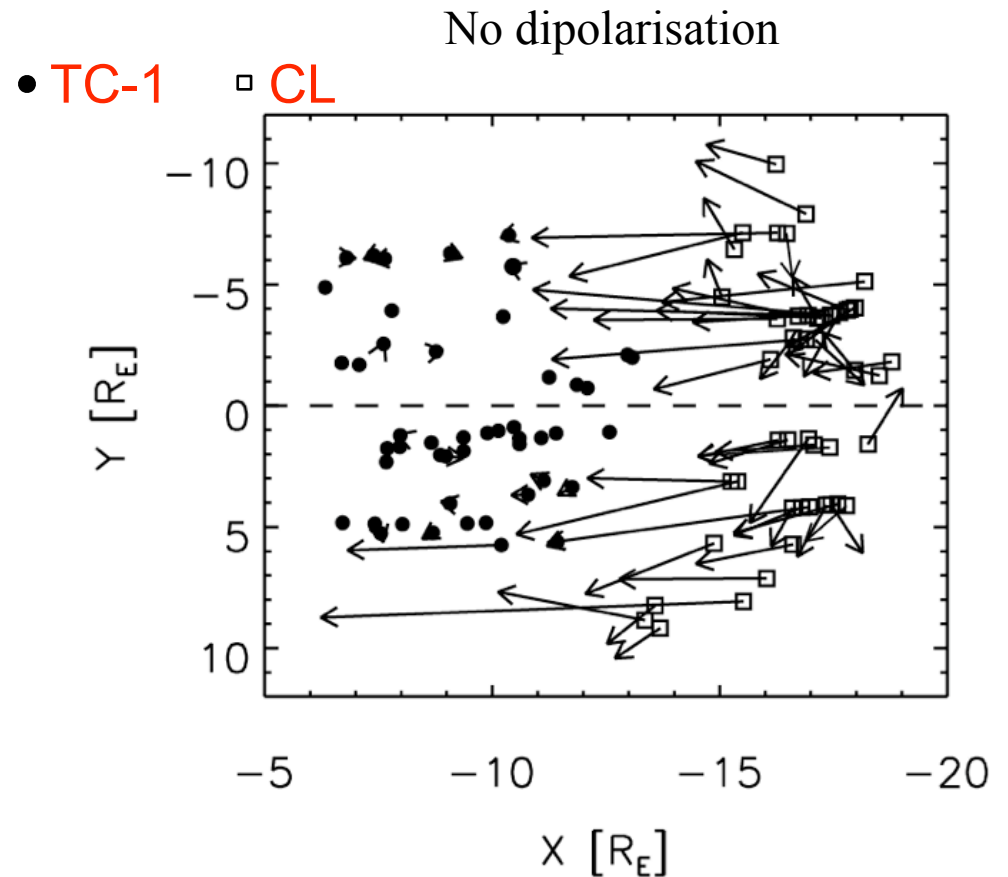
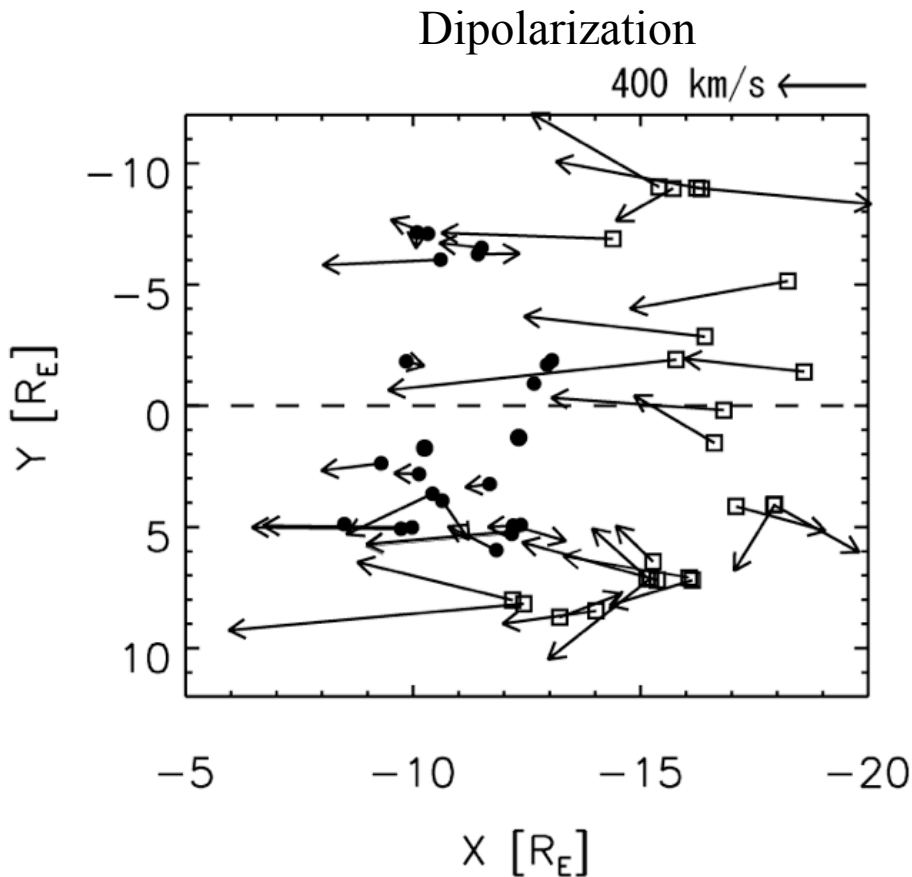


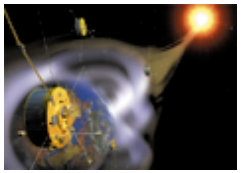
CLUSTER



# Bursty Bulk Flows and Dipolarization

- TC1 dipolarization events are far from  $X \sim -8$  RE
- The flow pattern observed by Cluster has similar features independent of the dipolarization at TC1

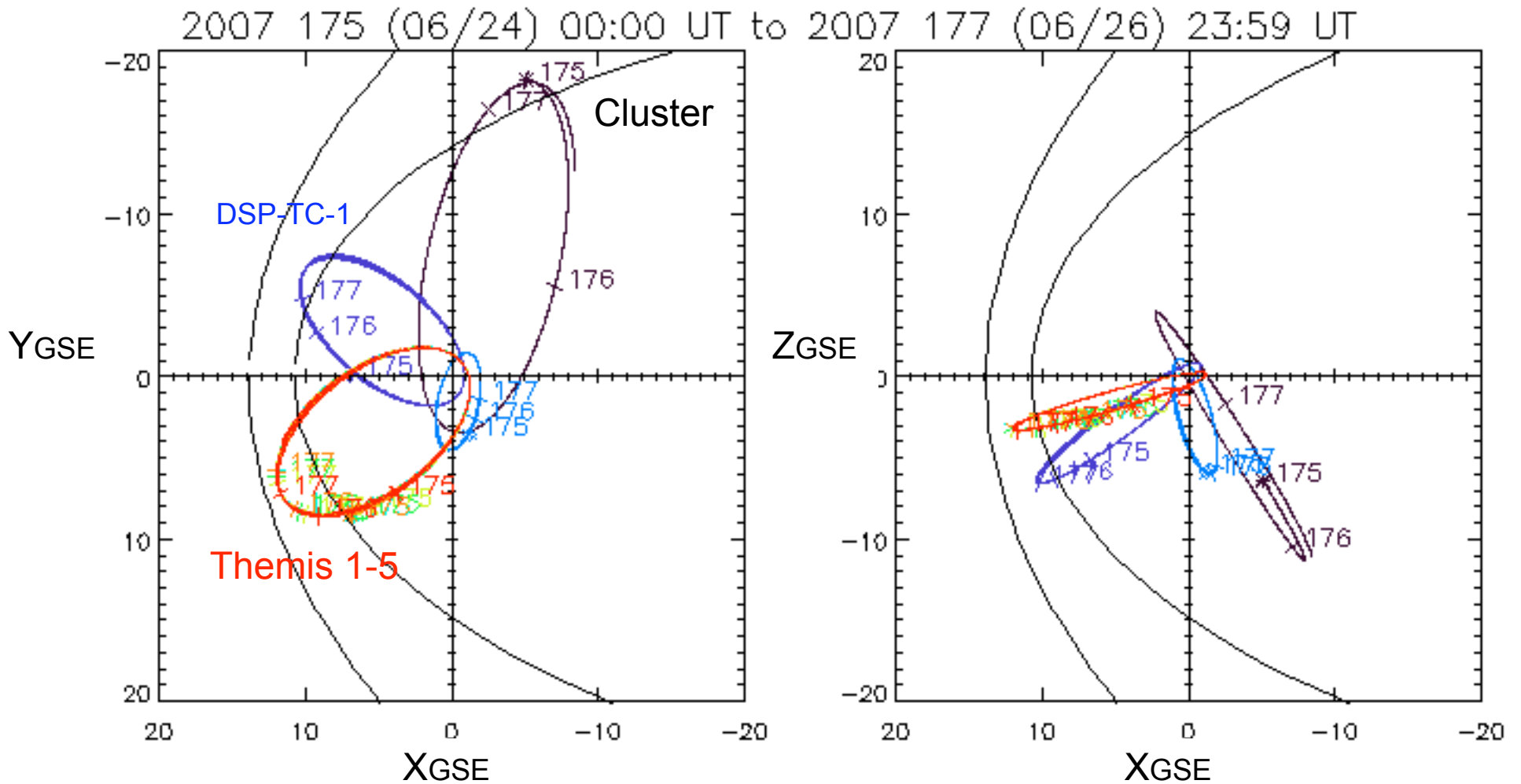




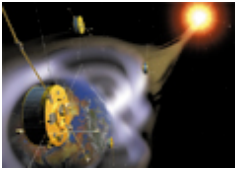
CLUSTER



# Conjunctions THEMIS-Cluster-Double Star, June 07 (before THEMIS full deployment)



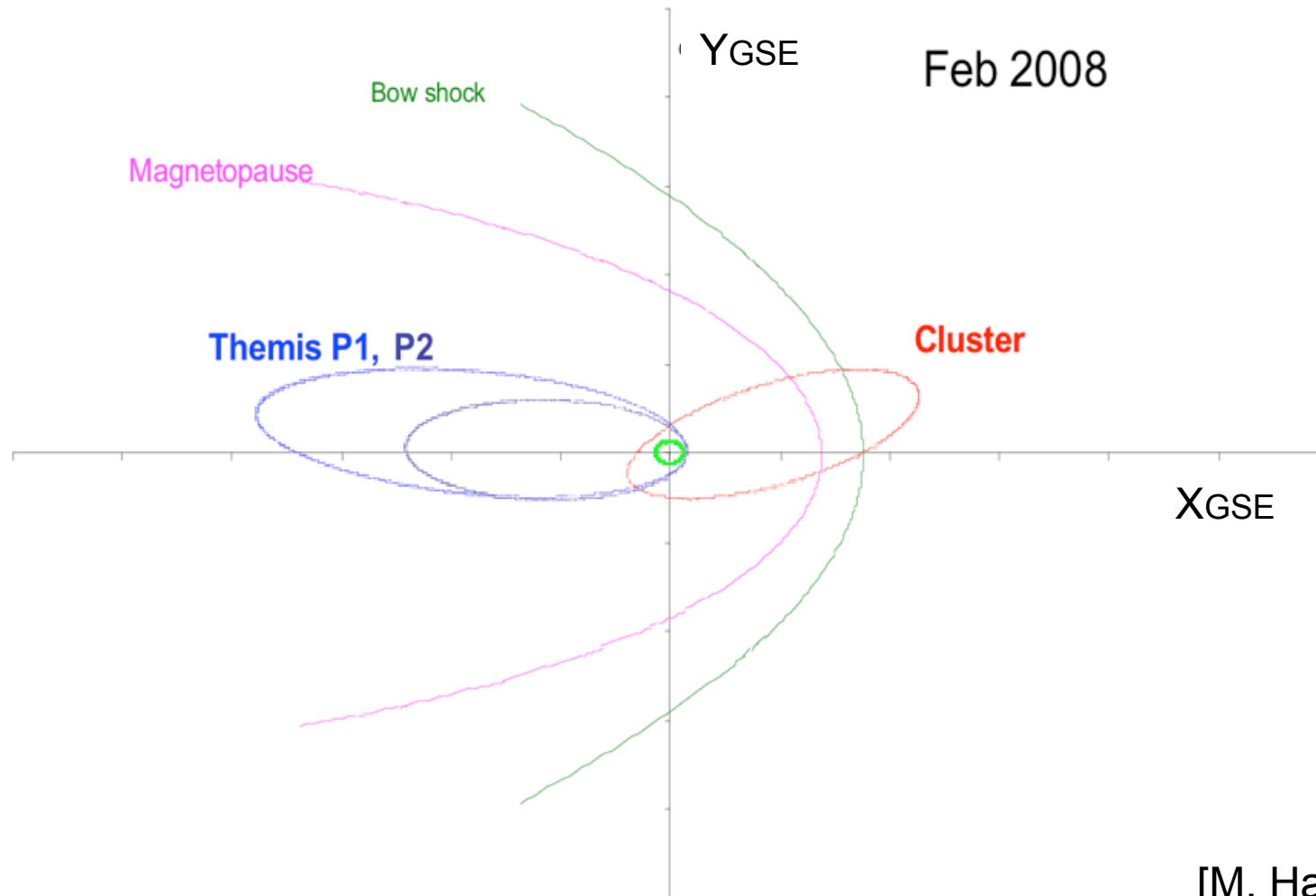
Note Double Star TC-1 re-entry in atmosphere beginning of Oct 2007



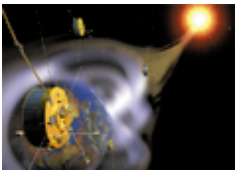
CLUSTER



## Conjunctions THEMIS-Cluster: tail and mid-altitude auroral zone



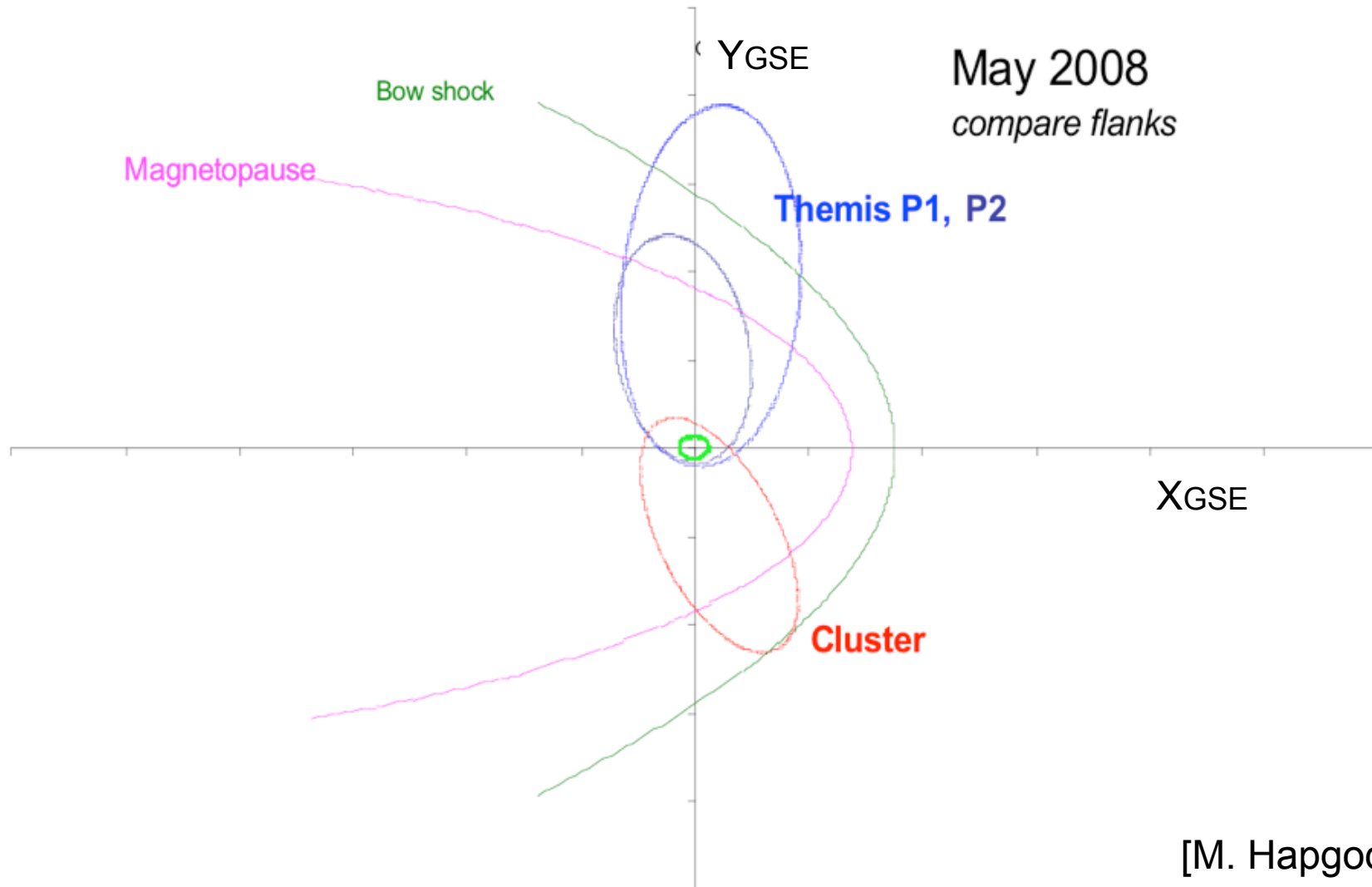
Note in 2009 Cluster will cross the auroral acceleration region around perigee

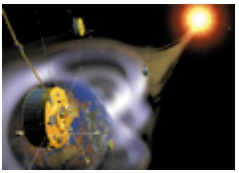


CLUSTER



## Conjunctions THEMIS-Cluster: compare dawn and dusk flanks

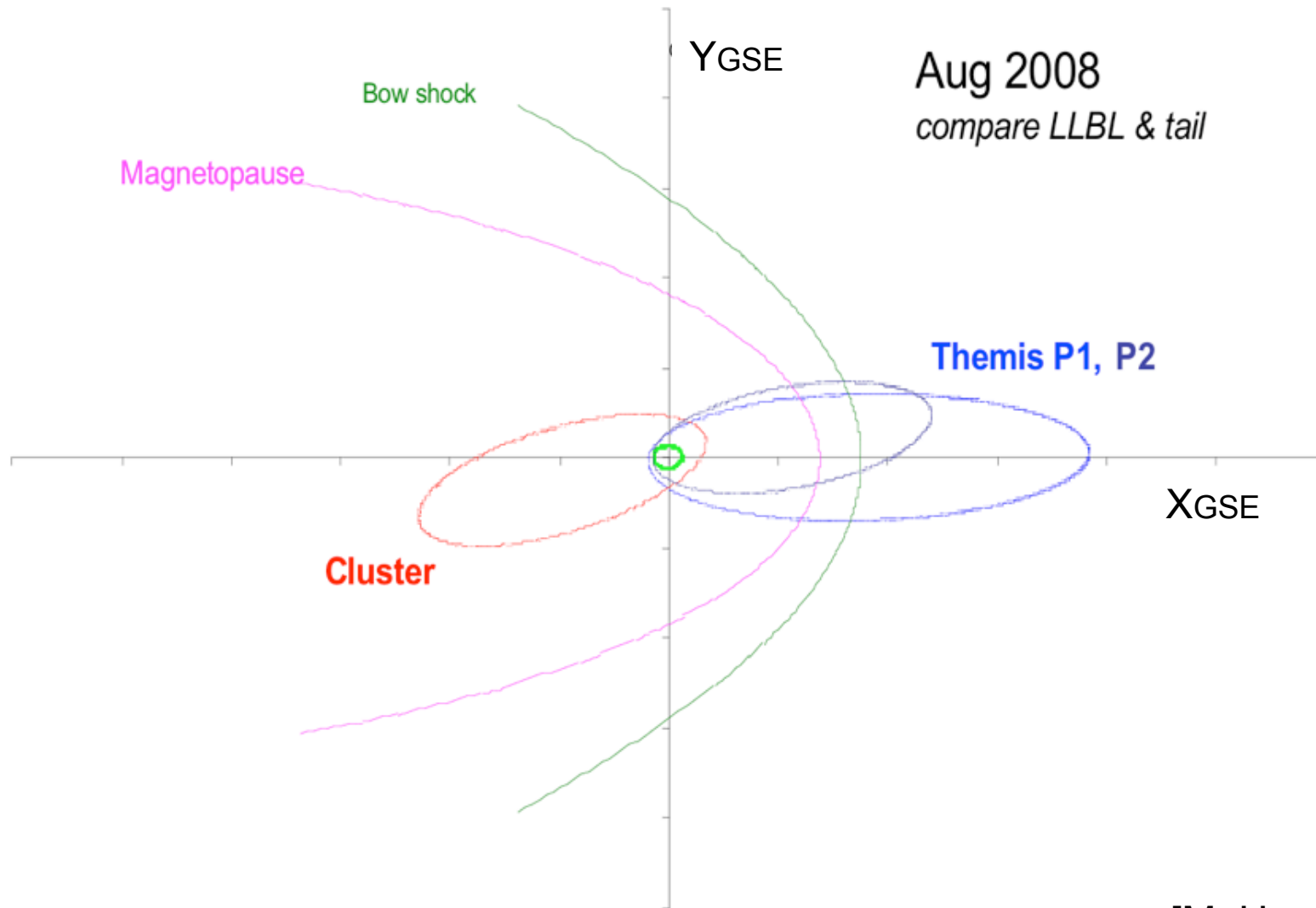




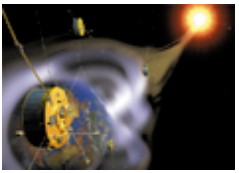
CLUSTER



# Conjunctions THEMIS–Cluster: LLBL-Tail & mid-alt. cusp-magnetopause



[M. Hapgood]



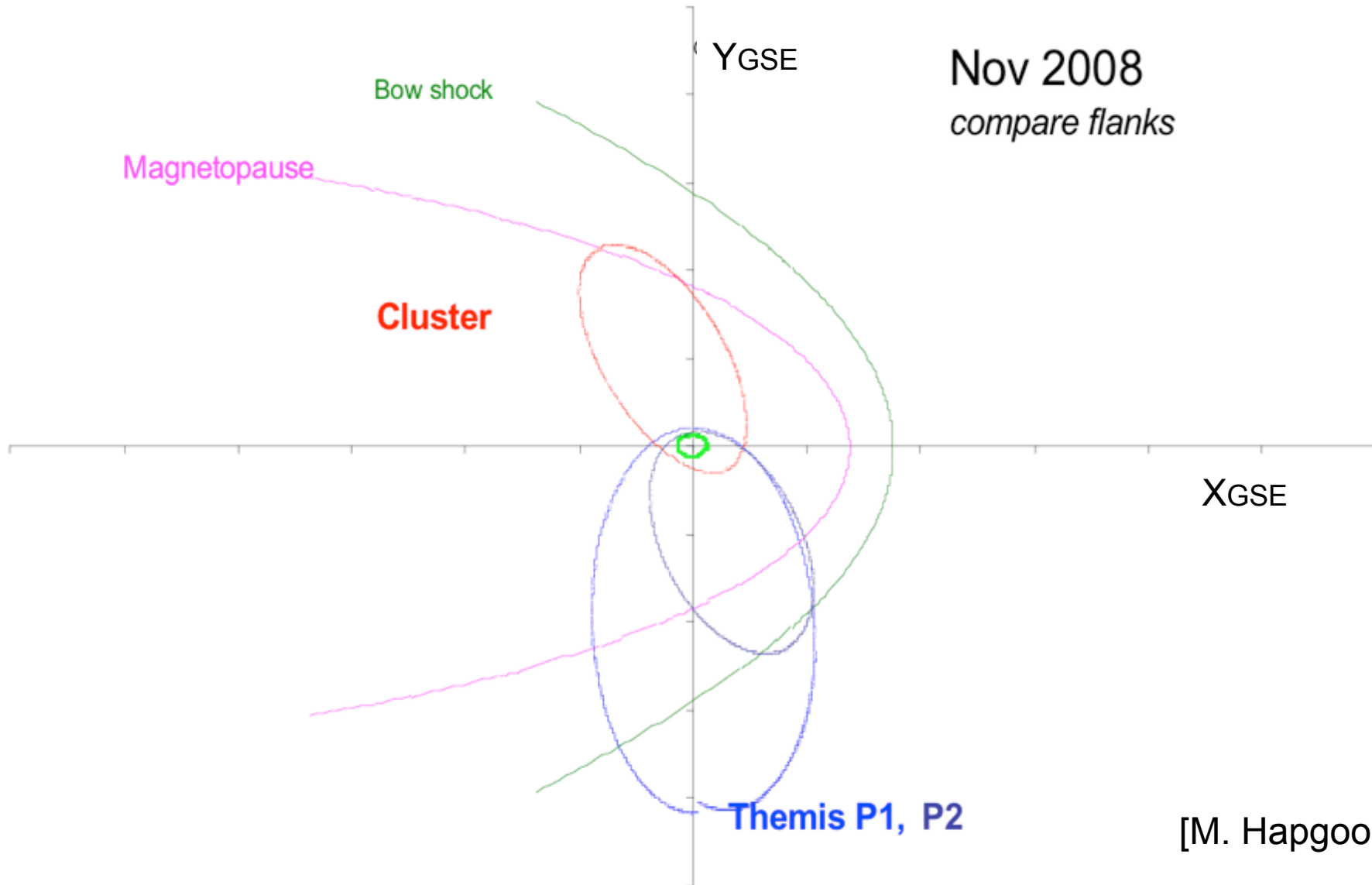
CLUSTER

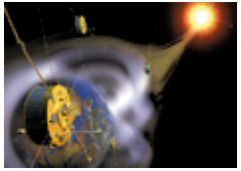


esa



# Conjunctions THEMIS–Cluster: compare dawn and dusk flanks





CLUSTER



# Summary and Conclusion

- Exciting new science with 11 spacecraft in the magnetosphere simultaneously: 5 THEMIS, 4 Cluster, and 2 Double Star
- Double Star TC-1 will cease operation in October 2007. TC-2 might continue if there are no thermal problems this summer
- Cluster has been extended to the end of 2009 (with review at end of 2007)