EVE and RHESSI

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Outline

- 0) SOL2010-06-12T00:57 as seen by HMI
- 1) What is EVE?
- 2) How good is EVE?
- 3) What can we learn from EVE and RHESSI?

SOL2010-06-12T00:57 (HMI)









ESP fluxes for SOL2010-06-12

EVE 304A for SOL2010-06-12

EVE 304A for SOL2010-06-12

EVE 304 Doppler sensitivity

EVE 304A for SOL2010-06-12

EVE 304A for SOL2010-06-12

Table 1: 30.4 nm Line Fits

Flare Phase	Time range	Irradiance	Centroid	Width	Redshift
		mW/m^2	nm	nm	$\rm km/s$
Preflare	00:50:09-00:54:29	4.26	30.3751	0.0312	
Impulsive	00:55:59-00:56:49	4.45	30.3754	0.0313	
Gradual	00:56:59-00:58:09	4.40	30.3753	0.0312	
Impulsive excess		0.197	30.3801	0.0322	48.88 ± 2.49^{a}
Gradual excess		0.137	30.3784	0.0305	32.05 ± 5.33^{a}
Impulsive-gradual					$16.8 {\pm} 5.9^{a}$

He 304 Charge Exchange

Peter et al. 1991

SOL2010-06-12 Charge Exchange

SOL2010-06-12 Charge Exchange

SOL2011-02-15 Fe XXIV

SOL2011-02-15 Fe XXIV

What can we learn from EVE and RHESSI?

- UV/EUV energy distribution for WLF/HXR
- Charge-exchange observations of α particles at .01-1 MeV/nucleon
- Transition-region explanation of 511 keV line width
- Flare dynamics, with big doses of imaging and modeling

Electrodynamic mapping of the solar interior to the corona

- Collisionality drops precipitously
- Hydrogen ionizes
- Plasma beta plummets
- Structurally important radiation escapes
- Vertical currents threading the photosphere carry the non-potential energy

Recall debate between Parker and Melrose in ApJ (1996)